

LETTER TO THE EDITOR

The *Gaia*-ESO Survey: Low- α element stars in the Galactic bulge[★]

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ABSTRACT

We take advantage of the *Gaia*-ESO Survey iDR4 bulge data to search for abundance anomalies that could shed light on the composite nature of the Milky Way bulge. The α -element (Mg, Si, and whenever available, Ca) abundances, and their trends with Fe abundances have been analysed for a total of 776 bulge stars. In addition, the aluminum abundances and their ratio to Fe and Mg have also been examined. Our analysis reveals the existence of low- α element abundance stars with respect to the standard bulge sequence in the $[\alpha/\text{Fe}]$ versus $[\text{Fe}/\text{H}]$ plane. Eighteen objects present deviations in $[\alpha/\text{Fe}]$ ranging from 2.1 to 5.3σ with respect to the median standard value. Those stars do not show Mg-Al anti-correlation patterns. Incidentally, this sign of the existence of multiple stellar populations is reported firmly for the first time for the bulge globular cluster NGC 6522. The identified low- α abundance stars have chemical patterns that are compatible with those of the thin disc. Their link with massive dwarf galaxies accretion seems unlikely, as larger deviations in α abundance and Al would be expected. The vision of a bulge composite nature and a complex formation process is reinforced by our results. The approach used, which is a multi-method and model-driven analysis of high resolution data, seems crucial to reveal this complexity.

Key words. Galaxy: bulge – Galaxy: abundances – Galaxy: stellar content

1. Introduction

The study of the Galactic bulge is rapidly unveiling its complexity. Because of its physical properties (metallicity distribution, age, spatial location, kinematical features, etc.), the bulge is at the crossroads of the other main Galactic components such the halo, thick disc and thin disc. As a consequence, the formation scenarios currently invoked are directly

linked to the general evolution of the Milky Way and to one crucial open question: What is the importance of fast versus secular evolution? Three main scenarios are proposed: i) in situ formation via dissipative collapse of a protogalactic gas cloud (Eggen et al. 1962); ii) accretion of substructures in a Λ CDM context (e.g. Scannapieco & Tissera 2003); and iii) secular formation from disc material through bar formation, vertical instability, buckling, and fattening, producing a boxy/peanut bulge (e.g. Martinez-Valpuesta & Gerhard 2013; Di Matteo et al. 2014).

Our knowledge of the bulge has improved relevantly in the recent years, however, many key open questions remain. By way of example, the number of bulge components, estimated from the metallicity distribution function, kinematical and structural data, is still under debate. Although up to five components have been suggested (e.g. Ness et al. 2013), recent studies (e.g. Rojas-Arriagada et al. 2014; Schultheis et al. 2017; Zoccali et al. 2017) show that only two components seem

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to be necessary. In this context, precise individual chemical abundances are crucial to disentangle the different evolutionary pathways responsible for the bulge formation. The first studies in this sense regarded mainly (but not exclusively) α -element abundances (e.g. McWilliam & Rich 1994; Bensby et al. 2013; Gonzalez et al. 2015) which, combined with iron abundances, can unveil important information regarding the initial mass function and the star formation history of a stellar system. Up to now, different analyses have revealed a single sequence in the $[\alpha/\text{Fe}]$ versus $[\text{Fe}/\text{H}]$ plane, flattening at metallicities lower than -0.37 ± 0.09 dex (e.g. Rojas-Arriagada et al. 2017), the metallicity at which the maximum of supernovae type Ia rate occurred. However, recent observations extending the analysis to other elements have already detected departures from what seemed to be a simple chemical evolutionary path, such as the existence of nitrogen overabundant stars (Schiavon et al. 2017b).

The goal of this paper is to take advantage of the *Gaia*-ESO Survey (GES; Gilmore et al. 2012; Randich et al. 2013) bulge data to search for abundance anomalies that could reveal the bulge composite nature. The data are described in Sect. 2, our results are presented in Sect. 3, and their interpretation is developed in the final Sect. 4.

2. *Gaia*-ESO Survey stellar parameters, distances and orbits

We have used the atmospheric parameters and individual chemical abundances (Fe, Mg, Si and Al) of bulge stars observed by GES, using the GIRAFFE spectrograph, and included in its fourth internal data release. The data are described in detail in Rojas-Arriagada et al. (2017). The initial total sample comprises 2320 red clump stars in 11 bulge fields, sampling the area $-10^\circ \leq l \leq +10^\circ$ and $10^\circ \leq b \leq -4^\circ$.

The HR21 set-up was employed for all the bulge stars. In addition, 172 stars in Baade's Window were observed with the HR10 set-up. Those stars are in common with the analysis of Hill et al. (2011).

As the search for abundance anomalies has to rely on very precise results, we applied the following selection criteria to the initial sample: $2.0 < \log g < 3.0$ dex, $\sigma_{T_{\text{eff}}} < 250$ K, $\sigma_{\log g} < 0.6$ dex, $\sigma_{[\text{Fe}/\text{H}]} < 0.1$ dex, $\sigma_{[\text{Mg}/\text{H}]} < 0.1$ dex, $\sigma_{[\text{Si}/\text{H}]} < 0.1$ dex, and $\sigma_{[\text{Al}/\text{H}]} < 0.1$ dex. This reduces the sample to 776 stars with very high quality parameters; 42 of these stars are observed with two set-ups. The mean signal-to-noise of the working sample is 280 ± 70 (deviation estimated from the MAD). We make use of the Rojas-Arriagada et al. (2017) spectroscopic distances, estimated for the same data set, to ensure that our selection is confined to the bulge region ($R_{\text{GC}} < 3.5$ kpc). The median error of the distances is 1.2 kpc (around 16%). Finally, OGLE proper motions available for the subsample of stars in Baade's Window (Sumi et al. 2004) allowed the estimation of the stellar orbital parameters. To this purpose, we adopted Model 4 presented in Fernández-Trincado et al. (2016) and composed of nonaxisymmetric potentials.

3. Search for α -element abundance anomalies

The main challenge of testing the existence of stars with non-standard low- α element abundances is demonstrating that they do not belong to the high error queue of a normal abundance distribution. To this purpose, we first carried out a fiducial median profile and a 1σ dispersion band, over 11 iron abundance bins for the $[\text{Mg}/\text{Fe}]$, $[\text{Si}/\text{Fe}]$, and $[\alpha/\text{Fe}]$ abundances. Secondly, we

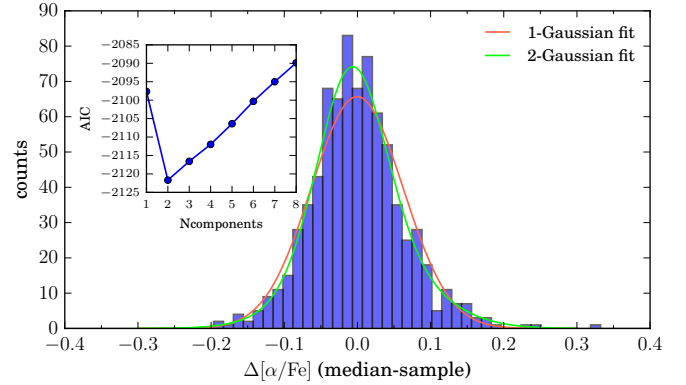


Fig. 1. Gaussian mixture method analysis of the $[\alpha/\text{Fe}]$ abundance distribution with respect to the median values along the sequence. The inner panel shows the Akaike information criterion for different numbers of components. The best fit (green curve) corresponds to the two-component mixture. The red curve shows the poorer quality fit of the single Gaussian component, whose probability of explaining the distribution is 21 times lower than the double Gaussian component.

computed, for each star, the difference in $[\text{Mg}/\text{Fe}]$, $[\text{Si}/\text{Fe}]$, and $[\alpha/\text{Fe}]$ with respect to the median values of the corresponding iron bin (in the sense median minus sample), called $\Delta[\text{Mg}/\text{Fe}]$, $\Delta[\text{Si}/\text{Fe}]$, and $\Delta[\alpha/\text{Fe}]$, respectively.

If there are low- α element stars in the bulge, the distribution of the $\Delta[\alpha/\text{Fe}]$ around the standard value should be 1) skewed and 2) not corresponding to a single Gaussian. We first applied a D'Agostino skewness test to the $\Delta[\alpha/\text{Fe}]$ abundance distribution. The skewness value of the $\Delta[\alpha/\text{Fe}]$ is 47.5 (p -value of $4.88\text{e-}11$), confirming the asymmetry. We tested the robustness of this result to the existence of strong outliers by reapplying the D'Agostino test considering only the stars within 2 sigma from the median $[\alpha/\text{Fe}]$ abundance value. Both the z-score of the test (6.9) and its p -value (0.03) seem to confirm that even the core of the distribution deviates from normality and it is skewed towards the low-alpha side. In addition, we applied a Gaussian mixture method (GMM; Rojas-Arriagada et al. 2016) to estimate the number of components of the distribution. The GMM algorithm constructs a generative model that consists in the specific Gaussian mixture that better predicts the data structure. We adopted the Akaike information criterion (AIC) as a cost function to assess the relative fitting quality between different proposed mixtures. The results of this analysis are shown in Fig. 1. Clearly, the single component solution can be excluded to explain the data distribution, highlighting the existence of an asymmetry.

Furthermore, the dispersion of the measured $[\alpha/\text{Fe}]$ abundances in metal-rich globular cluster giants can help us to realistically evaluate the errors on the bulge abundances. We considered two observed clusters within the same constraints as in our bulge analysis sample (cf. Sect. 2), i.e. NGC 104 ($[\text{Fe}/\text{H}] = -0.66 \pm 0.05$ dex, 30 stars) and NGC 5927 ($[\text{Fe}/\text{H}] = -0.33 \pm 0.06$ dex, 56 stars). The dispersion of the two clusters in the $[\alpha/\text{Fe}]$ abundances (0.035 dex and 0.045 dex, respectively), $[\text{Mg}/\text{Fe}]$ (0.04 and 0.05 dex, respectively) and $[\text{Si}/\text{Fe}]$ (0.05 and 0.06 dex, respectively) is clearly lower than those of the bulge distribution (0.06 dex for $[\alpha/\text{Fe}]$, 0.07 dex for $[\text{Mg}/\text{Fe}]$, and 0.08 dex for $[\text{Si}/\text{Fe}]$), reinforcing the hypothesis of a complex $[\alpha/\text{Fe}]$ versus $[\text{Fe}/\text{H}]$ sequence in the bulge.

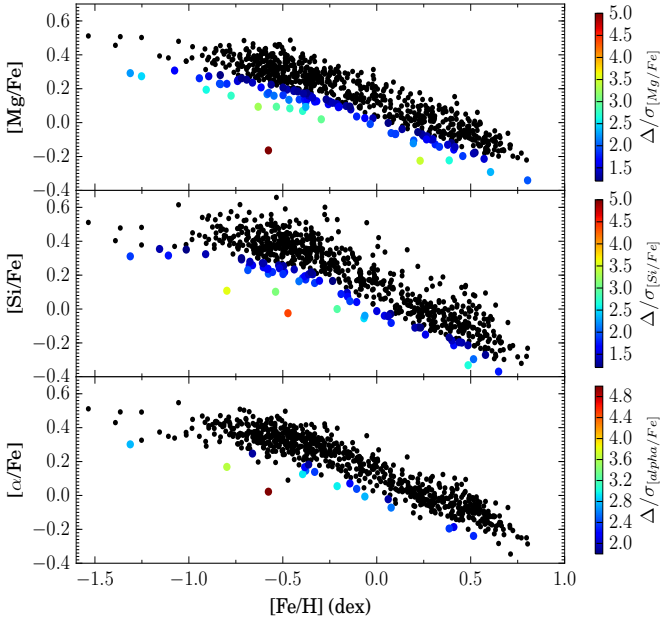


Fig. 2. Bulge distribution of $[Mg/Fe]$, $[Si/Fe]$, and $[\alpha/Fe]$ with respect to $[Fe/H]$ (upper, middle, and lower panels, respectively). The colour code shows the difference with respect to the median value ($\Delta[Mg/Fe]$, $\Delta[Si/Fe]$, and $\Delta[\alpha/Fe]$) expressed in units of standard deviation. Both Δ and σ are evaluated as a function of $[Fe/H]$.

3.1. Selection of stars separately from different α -elements.

First of all, we have analysed the abundances of Mg and Si and their trends with the Fe abundance. Figure 2 shows the $[Mg/Fe]$ versus $[Fe/H]$ distribution (upper panel), $[Si/Fe]$ versus $[Fe/H]$ distribution (middle panel), and $[\alpha/Fe]$ versus $[Fe/H]$ distribution defined as $([Mg/Fe] + [Si/Fe])/2$ (lower panel). Then, we identified the stars showing deviations larger than 1.2σ in $\Delta[Mg/Fe]$ and in $\Delta[Si/Fe]$ (colour coded in the upper and middle panels of Fig. 2). To ensure that the $[\alpha/Fe]$ anomaly is present in more than one element (avoiding possible Mg depleted globular cluster *escapées*; e.g. Carretta et al. 2009), we restricted the selection to the intersection of the low-Mg and low-Si subsamples. Finally, an additional condition was imposed: only the objects showing deviations greater than 2σ in $\Delta[\alpha/Fe]$ were selected. The final subsample of stars are colour coded in the bottom panel of Fig. 2. The subsample includes 18 objects with $[\alpha/Fe]$ underabundances ranging from 2.1 to 5.3σ . Of these, 90% are within 2.5 kpc from the Galactic centre and two-thirds are within 1.8 kpc.

A number of checks were performed to further test the reliability of the abundance anomalies. First of all, it was verified that the measurement dispersions in the parameters and abundances provided by the different GES analysis nodes (and taken as a proxy of the error) are similar for the anomalous stars and for the reference sample of 776 objects. In addition, the spectra of representative stars showing anomalies were compared to those of stellar twins (with maximum differences of 50 K for T_{eff} , 0.02 dex for $\log g$, and 0.03 dex for $[Fe/H]$). This is possible thanks to the fact that the bulge GES targets are restricted to the red clump and the chances of finding stellar twins in the sample are high. These comparisons allowed us to confirm that the Mg and Si abundances of the identified anomalous stars are in fact underabundant with respect to their twins of standard composition. As an example, the spectrum of the anomalous star

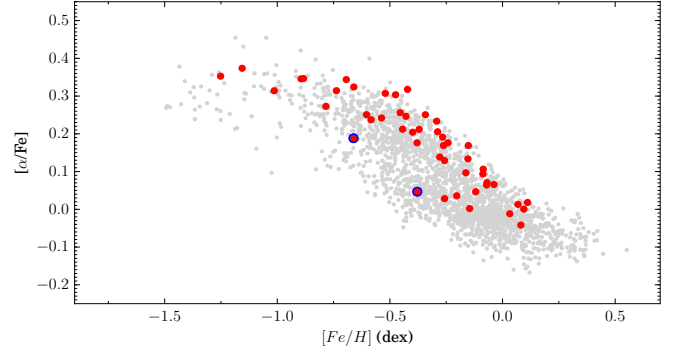


Fig. 3. Distribution of $[\alpha/Fe]$ vs. $[Fe/H]$ for the 48 stars of Baade's Window with high quality estimations of Mg, Si, and Ca abundances (red points). The two stars highlighted with blue circles are part of the group of stars identified to have Mg and Si underabundances in Fig. 2. Grey points show the $[\alpha/Fe]$ vs. $[Fe/H]$ distribution of GES disc stars.

17553025-4106299 (with $[Mg/Fe] = -0.16 \pm 0.08$ dex, $[Si/Fe] = 0.21 \pm 0.08$ dex) differs clearly in the α -element lines from that of its twin 18262565-3151577, having $[Mg/Fe] = 0.31 \pm 0.09$ dex and $[Si/Fe] = 0.37 \pm 0.06$ dex. An additional illustration of a spectral twins comparison is given in Fig. B.1. Third, the differences between the spectroscopic and photometric T_{eff} values (derived following González Hernández & Bonifacio 2009) of the stars presenting $[\alpha/Fe]$ abundance anomalies were checked and compared to those of the stars with standard compositions. No particular problems that could indicate specific errors in the spectroscopic T_{eff} of the anomalous stars were found. Moreover, the sensitivity to possible $\log g$ errors of the Mg and Si lines was verified using synthetic spectra. Although the estimated Si abundances are slightly sensitive to $\log g$ uncertainties, this is not the case for the Mg abundances. In the same line, the variation of the Mg and Si abundances with the typical $[M/H]$ uncertainties are at least six times smaller than the detected anomalies. Finally, we checked that the stars presenting abundance anomalies do not belong to a unique GES bulge field and GIRAFFE exposure, which could reveal possible problems with the data reduction, such as sky subtraction residuals. In conclusion, all the above-mentioned verification tests confirm the real α -poor nature of the identified bulge stars.

3.2. Baade's Window data

Particular attention was paid to the stars in Baade's Window (BW), for which both GES HR21 and HR10 spectra are available. Those stars with more reliable parameters and abundance estimations have measurements of an additional α element (calcium), and literature $[Mg/Fe]$ abundance determinations from Hill et al. (2011) for comparison. Figure 3 shows the distribution of $[\alpha/Fe]$ (defined as $[Mg + Si + Ca/Fe]$) versus $[Fe/H]$ for the 48 stars of BW with high quality estimations of Mg, Si, and Ca abundances (red points). The two red circles highlighted with blue edges correspond to the stars 18034317-3006349 ($[Fe/H] = -0.38 \pm 0.03$ dex) and 18032412-3003001 ($[Fe/H] = -0.66 \pm 0.02$ dex), which are already identified to have Mg and Si underabundances in Fig. 1. These two stars are confirmed to be $[\alpha/Fe]$ underabundant, even when Ca is considered. In addition, the star 18034317-3006349 seems to lay on the locus of the corresponding thin disc sequence, which is well below the standard $[\alpha/Fe]$ values of the bulge at its metallicity interval.

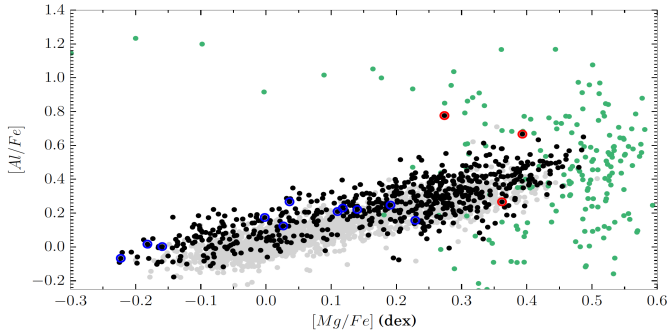


Fig. 4. [Al/Fe] abundances with respect to the [Mg/Fe] ratio for bulge stars (black points). Low- α stars and NGC 6522 stars are highlighted with blue and red circles, respectively. Disc stars and globular cluster values from Carretta et al. (2009) are shown as grey and green points.

Literature values from Hill et al. (2011) are available for these two stars. First, the GES iron abundances are compatible within the errors with the Hill et al. (2011) abundances. Second, the [Mg/Fe] ratio of Hill et al. (2011) for the star 18032412-3003001 (0.09 ± 0.18 dex) is also 1σ below the corresponding median [Mg/Fe] value for that metallicity interval. These authors do not report Mg abundance estimations for the other star.

Lastly, the orbits of the two stars seem confined to the bulge region. The derived radial apocentric distances for 18034317-3006349 and 18032412-3003001 are 2.75 ± 1.1 kpc and 1.87 ± 1.3 kpc, respectively. The corresponding maximum vertical amplitudes are 1.16 ± 0.2 kpc and 1.66 ± 0.4 kpc.

3.3. Aluminum abundances

Two aluminum lines in the HR21 spectra allowed the derivation of Al abundances for the GES bulge data. Figure 4 shows the [Al/Fe] abundance with respect to the [Mg/Fe] abundance for the studied bulge stars (black points). Low- α stars of Sect. 3.1 are highlighted with blue circles. As in Fig. 3, the abundances of disc stars are also included for comparison as light grey points. Finally, green points show the Carretta et al. (2009) values for GC stars (cf. their Fig. 5). First, the results presented in Fig. 4 show that the identified low- α bulge stars do not present GC abundance patterns. Their Al abundance seems to be in agreement with the Mg abundance, laying on the same sequence as bulge stars of standard composition. Moreover, the disc and bulge sequences are overlaid in the same locus of the figure, as expected.

Incidentally, three members of the globular cluster NGC 6522 ([Fe/H] = -1.14 ± 0.10 dex) were observed in the GES BW fields (indicated with red circles in Fig. 4). One of these stars, 18033857-3002434, presents a clear Al overabundance with [Al/Fe] = 0.78 ± 0.03 dex and a Mg depletion with respect to the other two ([Mg/Fe] = 0.27 ± 0.03 dex), with no sign of Si underabundance ([Si/Fe] = 0.38 ± 0.07) or Ca underabundance ([Ca/Fe] = 0.40 ± 0.07). Figure B.2 shows the fit of the two Al lines at 8772.8 and 8773.8 Å for this star, confirming the reliability of the estimated Al abundance and the continuum normalisation. As a consequence, the three NGC 6522 members present clearly a Mg-Al anti-correlation, as seen in Fig. 4. This result confirms the previous abundance measurements of Barbuy et al. (2009) of eight potential NGC 6522 members, and it reinforces the existence of multiple populations in this old cluster of relatively low mass.

4. Discussion

The analysis of the *Gaia*-ESO survey bulge data reveals the existence of low- α abundance stars. It is necessary to invoke a complex bulge phylogenesis to explain the reported intrinsic spread of [α /Fe] abundances at a given metallicity. The results presented here are compatible with the existence of a (at least) bimodal distribution of [α /Fe]. They also present a more detailed vision of the [α /Fe] versus [Fe/H] bulge trends than previous works analysing large samples of stars (e.g. Ness et al. 2013), thanks to the GES high resolution and high signal-to-noise data. These results bring to light the presence of bona fide low- α stars in the bulge that were already visible, without have been noted, in several studies in the literature (e.g. García Pérez et al. 2013, and references therein).

Several scenarios can be invoked to interpret this situation. First, only massive dwarf galaxies, such as the Sagittarius dwarf spheroidal, could have benefited from an efficient enough chemical enrichment to reach the metallicity range concerned by our analysis. In those cases, we would expect deviations in the [α /Fe] versus [Fe/H] plane of at least 3σ or more with respect to the standard [α /Fe] values of the bulge (e.g. McWilliam et al. 2013). Therefore, very few objects of our study have chemical patterns that are compatible with this explanation.

Second, the identified low- α abundance stars have chemical patterns that are compatible with those of the thin disc. A disc contribution to the bulge formation (e.g. Di Matteo et al. 2015) would indeed leave its imprints in the chemical abundances, as they are compatible with the observed outliers from the standard bulge sequence. In this sense, a possible spread (or bimodality) of the stellar ages (Haywood et al. 2016) in the bulge would be in agreement with this scenario. Nevertheless, the [α /Fe] spread could be present even at metallicities as low as -0.5 dex, challenging this interpretation.

On the other hand, our finding of multiple stellar populations in the cluster NGC 6522, added to the Schiavon et al. (2017a) results for other clusters, reinforces the link between the globular clusters and the nitrogen abundance anomalies of field stars reported by Schiavon et al. (2017b).

Finally, we would like to point out that the search of non-standard objects benefits from the methodology implemented by GES. First, the use of a multi-method spectral analysis allows the reduction of inevitable methodological errors. Moreover, the simultaneous analysis of several elements with similar nucleosynthetic origins is important to avoid errors affecting the lines of a particular element. Lastly, model-driven methods such as those employed in this study are more appropriate for the search of unexpected objects than data-driven methods, which work by construction in a close environment of knowledge.

In summary, the bulge composite nature and, therefore, its complex formation process is reinforced by the present results with all the strength of a multi-method abundance analysis of high resolution and high quality data.

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Appendix A: List of stars with low- α abundances

The list of stars identified in Sect. 3 as having anomalously low- α abundances is included in Table A.1.

Table A.1. *Gaia*-ESO Survey identifier, T_{eff} , $\log g$, $[\text{Fe}/\text{H}]$, $[\text{Mg}/\text{Fe}]$, $[\text{Si}/\text{Fe}]$, and $\Delta/\sigma[\alpha/\text{Fe}]$ for the identified stars with $[\alpha/\text{Fe}]$ underabundance.

Cname	T_{eff}	$\log g$	$[\text{Fe}/\text{H}]$	$[\text{Mg}/\text{Fe}]$	$[\text{Si}/\text{Fe}]$	$[\text{Al}/\text{Fe}]$	$\Delta/\sigma[\alpha/\text{Fe}]$	Set-up configuration
18032412-3003001	4674 ± 36	2.39 ± 0.14	-0.66 ± 0.03	0.24 ± 0.02	0.26 ± 0.02	—	2.0	HR10, HR21
18034317-3006349	4819 ± 95	2.65 ± 0.35	-0.38 ± 0.02	0.12 ± 0.06	0.17 ± 0.05	0.23 ± 0.06	2.5	HR10, HR21
17534571-4105165	4354 ± 37	2.06 ± 0.24	-0.39 ± 0.02	0.07 ± 0.06	0.18 ± 0.02	—	2.9	HR21
17553025-4106299	4606 ± 168	2.70 ± 0.26	-0.58 ± 0.02	-0.16 ± 0.08	0.21 ± 0.08	—	5.3	HR21
17562638-4129329	4561 ± 53	2.45 ± 0.42	-0.33 ± 0.05	0.12 ± 0.08	0.15 ± 0.07	—	2.4	HR21
17571064-4145311	4959 ± 87	2.74 ± 0.50	-0.38 ± 0.08	0.14 ± 0.02	0.19 ± 0.07	0.22 ± 0.05	2.2	HR21
17571166-4128521	4553 ± 49	2.55 ± 0.28	-0.06 ± 0.05	0.02 ± 0.09	-0.04 ± 0.06	0.12 ± 0.02	2.8	HR21
17573538-4136125	4809 ± 72	2.48 ± 0.24	-0.80 ± 0.02	0.23 ± 0.02	0.11 ± 0.09	0.16 ± 0.03	3.6	HR21
17582234-3449464	4560 ± 150	2.51 ± 0.59	0.07 ± 0.04	-0.06 ± 0.08	-0.08 ± 0.02	—	2.6	HR21
18030869-2956389	4607 ± 105	2.78 ± 0.45	-0.11 ± 0.07	0.03 ± 0.09	0.04 ± 0.03	0.26 ± 0.04	2.4	HR21
18033088-3003563	4551 ± 43	2.71 ± 0.08	0.41 ± 0.02	-0.16 ± 0.05	-0.21 ± 0.02	0.00 ± 0.03	2.1	HR21
18035188-3005558	4574 ± 65	2.86 ± 0.28	-0.21 ± 0.08	0.11 ± 0.02	0.00 ± 0.09	0.21 ± 0.03	3.0	HR21
18233838-3413015	4625 ± 64	2.98 ± 0.29	0.38 ± 0.07	-0.22 ± 0.03	-0.17 ± 0.02	-0.07 ± 0.04	2.4	HR21
18244002-2449329	5000 ± 216	2.83 ± 0.52	-1.31 ± 0.08	0.29 ± 0.03	0.31 ± 0.02	—	2.8	HR21
18250175-2514127	4463 ± 91	2.79 ± 0.08	0.51 ± 0.02	-0.18 ± 0.08	-0.29 ± 0.04	0.02 ± 0.02	2.3	HR21
18251816-2529497	4663 ± 45	2.78 ± 0.37	-0.36 ± 0.07	0.16 ± 0.04	0.20 ± 0.03	—	2.0	HR21
18352178-2659169	4636 ± 57	2.85 ± 0.33	0.06 ± 0.03	0.00 ± 0.02	-0.04 ± 0.04	0.17 ± 0.06	2.0	HR21
18360308-2716364	4664 ± 36	2.43 ± 0.28	-0.14 ± 0.06	0.09 ± 0.03	0.04 ± 0.08	—	2.2	HR21

Notes. The reported uncertainties correspond to the measurement dispersion between the different analysis methods.

Appendix B: Visual inspection of spectral lines

Figure B.1 presents two examples of α -element lines (the CaII line at 8498.0 Å and the MgI line at 8806.7 Å) allowing the comparison of the spectra of two stellar twins. The abundances differences in Ca and Mg are reported in the figure.

Figure B.2 shows two aluminum lines of one Al-rich star of the globular cluster NGC 6522.

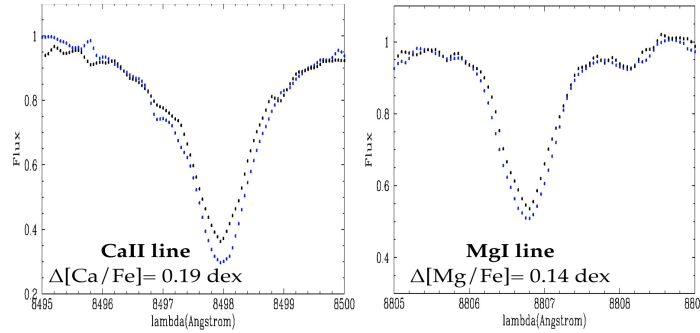


Fig. B.1. Comparison of the spectra of two stellar twins: the low- α star 18034317-3006349 (black points) and the star 18035553-3002562 with a standard $[\alpha/\text{Fe}]$ composition (blue points).

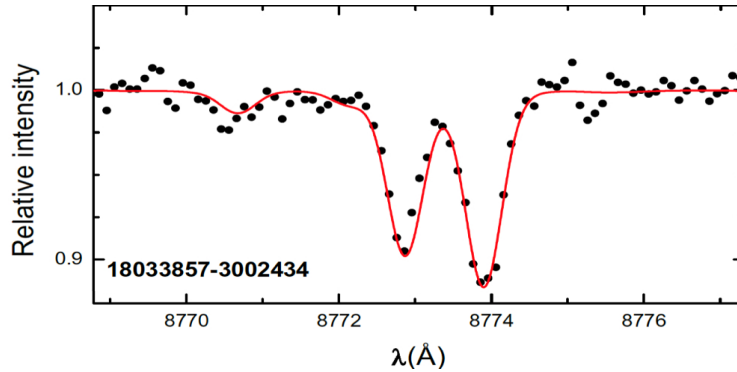


Fig. B.2. Spectrum of the star 18033857-3002434 (black points), around the two Al lines at 8772.8 and 8773.8 Å. The red line shows a fit with a synthetic spectrum corresponding to $[\text{Al}/\text{Fe}] = 0.78 \pm 0.03$ dex.